

Numerical GRMHD Simulations of Self-Gravitating

Collapsing Stars

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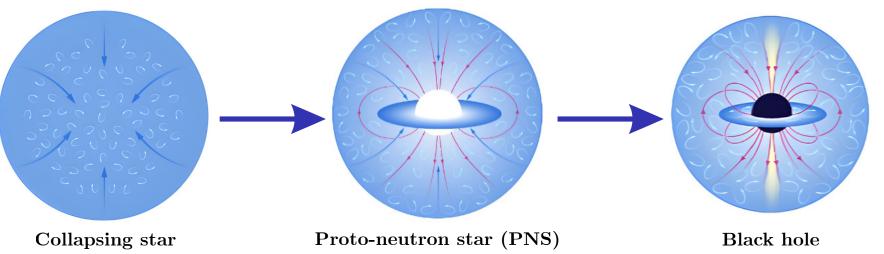
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ABSTRACT

- We present, for the first time, 3D GRMHD simulations of collapsars, including both the black hole gravity and the **self-gravity of the stellar envelope** in the framework of general relativity.
- Models with self-gravity (SG) and without self-gravity (NSG) are compared under identical initial conditions to assess the effects of self-gravity on the jet properties and the electromagnetic transient.
- This work is a **continuation** of the two-dimensional study on self-gravity effects in collapsars (Janiuk et al. 2023).

COLLAPSAR MODEL

- The collapse of a rapidly rotating massive star leads to long gamma-ray bursts (Woosley 1993).
- The inner core collapses into a rapidly rotating **Kerr black hole** (in some cases via a proto-neutron star stage).
- The **progenitor** must have **sufficient angular momentum** to form an equatorial **accretion disk**.
- Relativistic jets are launched along the rotation axis of the black hole.



METHODS AND SELF-GRAVITY IMPLEMENTATION

Figure 1: Schematic view of the collapse of a massive star to a black hole with the proto-neutron star stage (Gottlieb et al. 2024).

• We performed numerical simulations using the code **HARM** (Gammie et al. 2003; Noble et al. 2006), which solves the following GRMHD equations (the continuity equation, the conservation of energy and momentum, and the induction equation governing the evolution of the magnetic field):

$$\nabla_{\mu}(\rho u^{\mu}) = 0, \quad \nabla_{\mu}(T^{\mu}_{\nu}) = 0, \quad \nabla_{\mu}(u^{\mu}b^{\nu} - u^{\nu}b^{\mu}) = 0.$$

• The implementation of **self-gravity** is based on the **time-dependent Kerr metric** approximation:

Kerr–Schild metric (depends on time): simulations (does not depend on time): $g_{\mu\nu}(r,\theta,\phi,M_0,a_0)$

 $g_{\mu\nu}(r,\theta,\phi,M_0,a_0,\Delta M(t),\delta M(t,r),\Delta a(t),\delta a(t,r))$

Our implementation of the self-gravitating dynamical

• The motivation for including self-gravity is the large mass difference between the initial black hole and the stellar envelope $(3 M_{\odot} \text{ and } 25 M_{\odot})$

	$M_{ m BH} \ (M_{\odot})$	$M_{star} \ (M_{\odot})$	a_0	S	Magnetic Field	Perturbation (%)	Jet
	3	25	0.8	2	vertical	5	yes
	3	25	0.8	2	vertical	5	yes
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• Resolution was set to: $(N_r, N_\theta, N_\phi) = (384, 192, 128)$

The usual static Kerr–Schild metric used in GRMHD

$\Delta M(t) = \int_0^t \int_0^{2\pi} \int_0^{\pi} \sqrt{-g} \, T_t^r \, d\theta \, d\phi \, dt'$ $\delta M(t,r) = \int_{r_{\rm h}}^{r} \int_{0}^{2\pi} \int_{0}^{\pi} \sqrt{-g} \, T_{t}^{t} \, d\theta \, d\phi \, dr'$ $\Delta J(t) = \int_0^t \int_0^{2\pi} \int_0^{\pi} \sqrt{-g} \, T_\phi^r \, d\theta \, d\phi \, dt'$ $\delta J(t,r) = \int_{0}^{r} \int_{0}^{2\pi} \int_{0}^{\pi} \sqrt{-g} \, T_{\phi}^{t} \, d\theta \, d\phi \, dr'$

JET LAUNCHING AND PROPERTIES

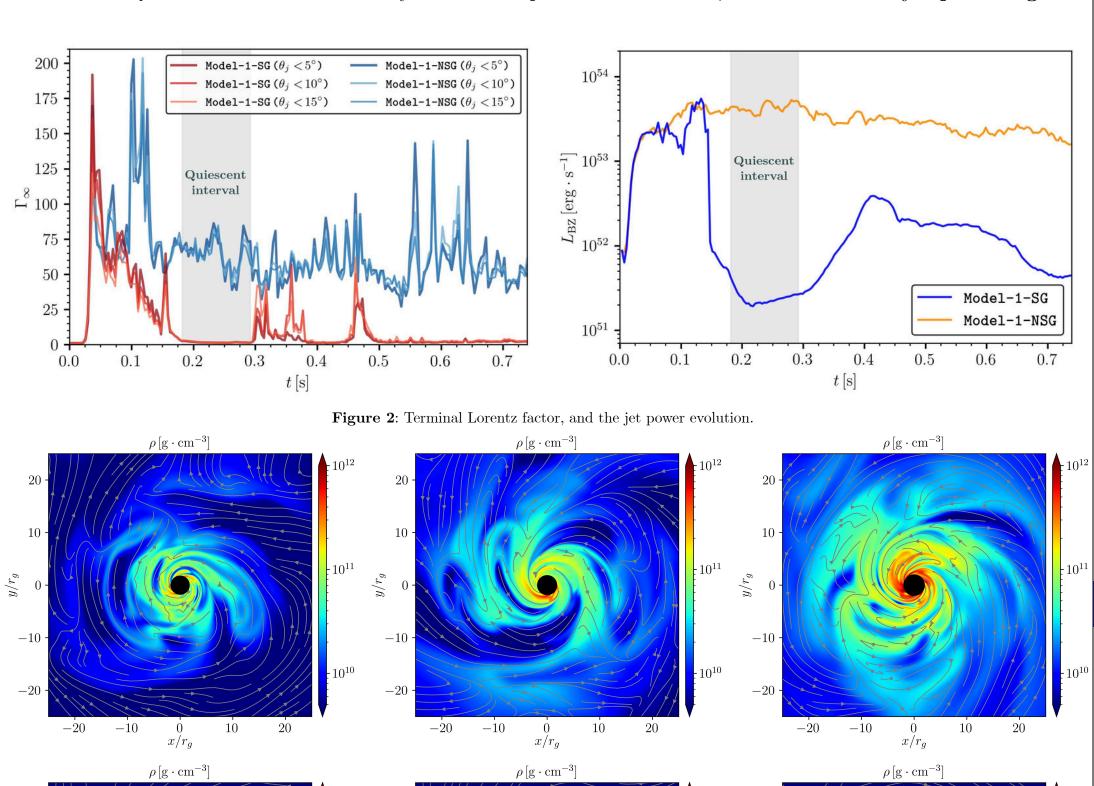
- The main mechanism responsible for launching Poynting jets is the Blandford-Znajek process (Blandford & Znajek 1977). The majority of the energy is initially stored in the electromagnetic field, then dissipated, converted into the kinetic energy of particles, and eventually partially emitted as highly energetic radiation.
- To measure the energetics of jets, we use the Lorentz factor estimated at infinity. This parameter is defined under the assumption that, at infinity, all forms of energy are converted into baryonic bulk kinetic energy.

$$\mu = \Gamma_{\infty} = -\frac{T_t^r}{\rho u^r}.$$

• Since, in Poynting jets, almost the entire energy is stored in the electromagnetic field, we calculate the **power of** the jet by taking the electromagnetic flux through the enclosed surface above the event horizon.

$$P_{jet} = -\int_{0}^{2\pi} \int_{0}^{\pi} \sqrt{-g} \left(b^{2} u^{r} u_{t} - b^{r} b_{t} \right) d\theta d\phi.$$

- The model with self-gravity exhibits more variable jet emission, with intervals of halted emission, which is generally more consistent with observed prompt gamma-ray emissions.
- The main Quiescent interval is directly related to rapid mass accretion, which leads to the jet quenching.



ACCRETION FLOW AND BLACK HOLE EVOLUTION

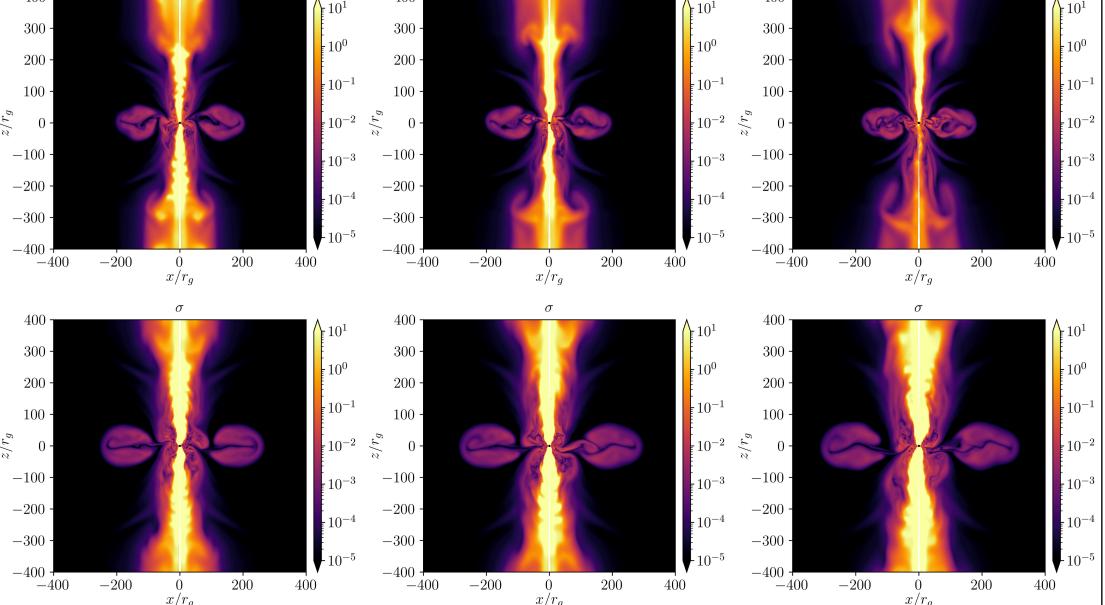


Figure 4: Poloidal maps of magnetization σ for Model-1-SG (top row) and Model-1-NSG (bottom row). Columns from left to right correspond to times t = 0.074 s, 0.092 s, and 0.111 s, respectively. The gradual jet quenching in the self-gravity model is responsible for the Quiescent interval.

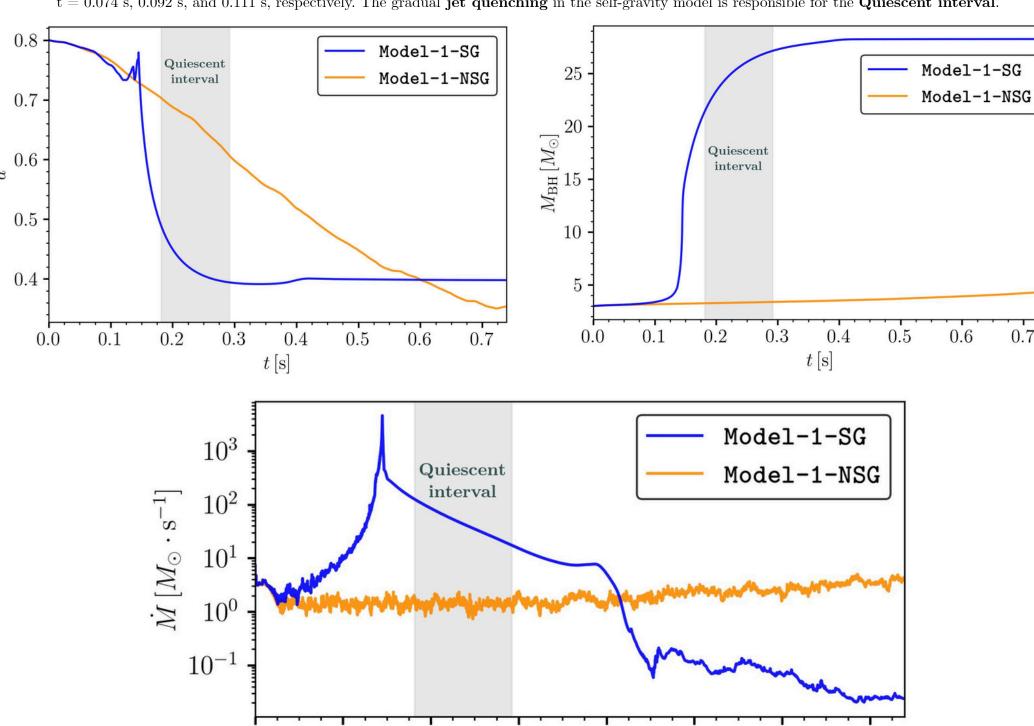


Figure 5: Evolution of the black hole spin, mass, accretion energy rate.

0.3

0.4

t[s]

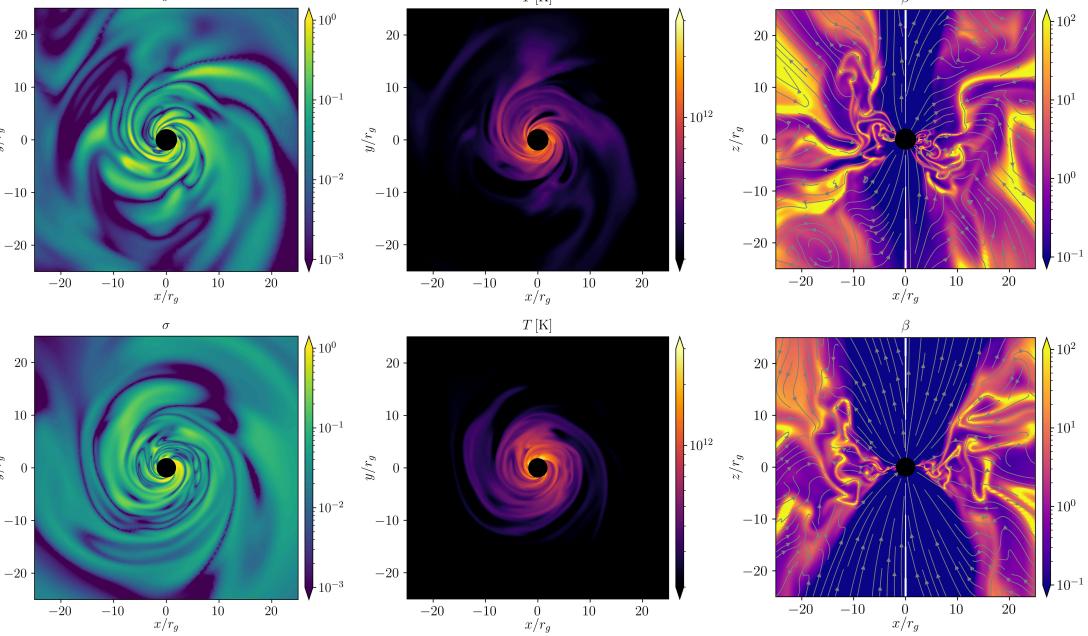


Figure 6: Maps of magnetization, temperature, and plasma beta for Model-1-SG (top row) and Model-1-NSG (bottom row). The columns correspond to t = 0.111 s. In the self-gravity model, due to the mass accumulation, the system becomes more chaotic and the mass begins to penetrate the jet funnel.

CONCLUSIONS

- Model-1-SG shows that the mass accumulation leads to the gradual suppression of jet emission, resulting in the Quiescence interval. Then, emission reappears, when when most of the stellar envelope mass is accreted.
- Self-gravity leads to a **faster evolution** of the **black hole spin and mass**. Additionally, we observe that the model without self-gravity extracts more rotational energy from the black hole.
- There is **no** substantial **difference** in the **jet launching start time** between the two models.
- The opening angle of the jet in Model-1-SG is smaller than in Model-1-NSG, which is related to additional inward pressure to the jet funnel in the model with self-gravity.

The main conclusion:

Some quiescent intervals in the prompt emission of long GRBs could be attributed to a mass accumulation in the vicinity of the black hole, which temporarily suppresses jet activity.

REFERENCES

Blandford & Znajek (1977). "Electromagnetic extraction of energy from Kerr black holes." In: Monthly Notices of the Royal Astronomical Society 179, p. 433–456. Gottlieb et al. (2024), "She's Got Her Mother's Hair: Unveiling the Origin of Black Hole Magnetic Fields through Stellar to Collapsar Simulations". In: The Astrophysical Journal Letters 976, p. 1 Gammie et al. (2003). "HARM: A Numerical Scheme for General Relativistic Magnetohydrody-namics". In: The Astrophysical Journal 589.1, p. 444-457. Janiuk et al. (2023). "Self-gravitating collapsing star and black hole spin-up in long gamma ray bursts". In: Astronomy & Astrophysics 677, A19.

Noble et al. (2006). "Primitive Variable Solvers for Conservative General Relativistic Magnetohydrodynamics". In: The Astrophysical Journal 641.1, p. 626.

Woosley (1993). "Gamma-Ray Bursts from Stellar Mass Accretion Disks around Black Holes". In: The Astrophysical Journal 405, p. 273.

Figure 3: Mass density, overlaid with magnetic field lines for Model-1-SG (top row) and Model-1-NSG (bottom row). Columns from left to right correspond to times t = 0.074 s, 0.092 s, and 0.111 s, respectively. In the model with self-gravity, we observe the mass accumulation in the vicinity of the black hole.

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